$\boxed{\textbf{G P21}}$ The puzzling Maia candidate star α Draconis

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Since the postulation of the enigmatic class of Maia variables in 1950 their existence is still disputed. If they indeed exist, they are located in the Hertzsprung–Russell diagram between the blue egde of the classical instability strip and the red border of the slowly pulsating B stars and consequently pulsation is not expected.

We revealed luminosity variations in two independent photometric time series of the A0 III Maia candidate star α DRACONIS with a period of about 53 minutes and an amplitude less than one millimagnitude. A spectroscopic time series results in radial velocity variations with about 40 ms^{-1} amplitude and the same period as determined photometrically. α DRACONIS is a single–lined spectroscopic binary system with an orbit period of about 50 days. The small distance of the two stars (about 0.4 AU) might cause pulsation driven by tidal forces. If this assumption is true, and new spectroscopic observations indicate this, the amplitude of the variations should be modulated with the orbit period. Further observations at various orbit phases are need to clarify this problem.